

On the evolution and environmental dependence of the star formation rate versus stellar mass relation since $z \sim 2$

Yusei Koyama,^{1,2*} Ian Smail,³ Jaron Kurk,⁴ James E. Geach,⁵ David Sobral,⁶ Tadayuki Kodama,^{2,7} Fumiaki Nakata,^{2,7} A. M. Swinbank,³ Philip N. Best,⁸ Masao Hayashi,² Ken-ichi Tadaki⁹

¹*Department of Physics, Durham University, South Road, Durham DH1 3LE, UK*

²*National Astronomical Observatory of Japan, Mitaka, Tokyo 181-8588, Japan*

³*Institute for Computational Cosmology, Durham University, South Road, Durham DH1 3LE, UK*

⁴*Max-Planck-Institut für Extraterrestrische Physik, Postfach 1312, Giessenbachstrasse, D-85741 Garching, Germany*

⁵*Department of Physics, McGill University, 3600 Rue University Montréal, Québec H3A 2T8, Canada*

⁶*Leiden Observatory, Leiden University, PO Box 9513, NL-2300 RA Leiden, the Netherlands*

⁷*Subaru Telescope, National Astronomical Observatory of Japan, 650 North A'ohoku Place, Hilo, HI 96720, USA*

⁸*SUPA, Institute for Astronomy, Royal Observatory of Edinburgh, Blackford Hill, Edinburgh EH9 3HJ*

⁹*Department of Astronomy, Graduate School of Science, The University of Tokyo, Tokyo 113-0033, Japan*

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ABSTRACT

This paper discusses the evolution of the correlation between galaxy star formation rates (SFRs) and stellar mass (M_*) over the last ~ 10 Gyrs, particularly focusing on its environmental dependence. We first present the mid-infrared (MIR) properties of the $H\alpha$ -selected galaxies in a rich cluster Cl0939+4713 at $z = 0.4$. We use wide-field Spitzer/MIPS24 μ m data to show that the optically red $H\alpha$ emitters, which are most prevalent in group-scale environments, tend to have higher SFRs and stronger dust extinction than the majority population of blue $H\alpha$ sources. With a MIR stacking analysis, we also find that the median SFR of $H\alpha$ emitters increases in higher-density environment at $z = 0.4$, and this trend is confirmed for both red and blue galaxies. The trend becomes much less significant when we compare their specific SFR (SSFR), although we find that there still remains a weak, but positive correlation between SSFR and galaxy number density. We then discuss the environmental dependence of the SFR versus M_* relation for star-forming galaxies since $z \sim 2$, by compiling comparable $H\alpha$ -selected galaxy samples in both distant cluster environments (from MAHALO-Subaru) and field environments (from HiZELS). Based on these large, narrow-band selected $H\alpha$ emitter samples, we find that the SFR– M_* relation does not show detectable environmental dependence since $z \sim 2$. Our data also suggest that the SSFR of $H\alpha$ -selected galaxies (at $\log(M_*/M_\odot) = 10$) evolves as $(1+z)^3$, and this evolutionary speed seems to be independent of the environment. Therefore we suggest that the primary star-formation quenching mechanism in cluster environments is a fast-acting process at any time in the history of the Universe since $z \sim 2$.

Key words: galaxies: clusters: individual: Cl0939+4713 — galaxies: evolution — large-scale structure of Universe.

1 INTRODUCTION

Galaxy formation and evolution is strongly dependent on environment. In the local Universe, galaxies in cluster environments are mostly passive (red), early-type galaxies (e.g. Dressler 1980; Goto et al. 2003), and there is a clear trend that the star formation activity of galaxies tends to be lower in high-density environ-

ment than low-density fields (e.g. Lewis et al. 2002; Gómez et al. 2003; Tanaka et al. 2004; Balogh et al. 2004). Therefore it is believed that the star formation activity of galaxies is affected by their surrounding environment during the course of the cluster or group assembly process. The environmental trends are also seen in the distant Universe (e.g. Kodama et al. 2001; Postman et al. 2005; Quadri et al. 2012), and some recent studies suggest that the star formation–density relation may be reversed at $z \gtrsim 1$ (Elbaz et al. 2007; Cooper et al. 2008; Ideue et al. 2009; Tran et al. 2010). This

* E-mail: yusei.koyama@durham.ac.uk

“reversal” of the star formation–density relation in the early Universe is still under debate (Patel et al. 2009; 2011), because the results could be highly uncertain depending on the sample definitions or selection method (e.g. Feruglio et al. 2010; Sobral et al. 2011; Popesso et al. 2011). Nevertheless, some recent studies which focus on individual galaxy clusters indeed find a hint that a fraction of galaxies in distant ($z \sim 1$) cluster outskirts or intermediate-density environments are showing boosted activity (e.g. Marcillac et al. 2007; Poggianti et al. 2008; 2009; Koyama et al. 2010; Geach et al. 2011), suggesting an accelerated galaxy evolution at the site of active cluster assembly.

A growing number of studies have revealed important roles of the group-scale environment or in-falling regions around rich clusters, by studying e.g. galaxy colours or morphologies (e.g. Kodama et al. 2001; Wilman et al. 2008; Balogh et al. 2011). In addition, optical emission-line surveys or MIR–FIR observations of distant clusters have brought some new insights into the (obscured) nature of star forming galaxies in cluster environment. A prominent example is *dusty red* galaxies, which are reported to populate the outskirts of rich galaxy clusters out to $z \sim 1$ (e.g. Koyama et al. 2008; 2010; see also Wolf, Gray, & Meisenheimer 2005; Geach et al. 2006; Verdugo, Ziegler, & Gerken 2008; Wolf et al. 2009; Tran et al. 2009). These studies suggest that such dusty red galaxies are a key population for understanding the physics of environmental effects.

In this respect, wide-field $H\alpha\lambda 6563$ emission-line surveys of distant galaxy clusters are a powerful method to pinpoint the location of this key population. This is not only because the $H\alpha$ line is less affected by dust extinction compared to star-formation indicators at rest-frame ultra-violet wavelengths (e.g. $[OII]\lambda 3727$), but also because emission line surveys with narrow-band (NB) filters allow us to effectively pick out galaxies from a narrow redshift slice. The latter is particularly important for cluster studies because the effect of contamination could be a concern. Koyama et al. (2011) performed a wide-field $H\alpha$ emission-line survey of a rich cluster, Cl0939+4713 ($z = 0.41$), and find a strong concentration of optically red star-forming galaxies in the group-scale environment around the cluster. While we argued in Koyama et al. (2011) that the excess of the red star-forming galaxies suggests an enhancement of dust-obscured star formation in the group environment, a firm conclusion still await a direct measurement of dust-enshrouded star-formation in those galaxies, because even $H\alpha$ lines are reported to be heavily extinguished in extremely dusty galaxies (e.g. Poggianti & Wu 2000; Geach et al. 2006; Koyama et al. 2010). Therefore, the first goal of this paper is to directly unveil the nature of this red star-forming population in distant group environments using mid-infrared (MIR) observations.

Another important parameter that drives galaxy evolution is the stellar mass (M_*) of galaxies. A correlation has been claimed between galaxy SFR and M_* for star-forming galaxies in the local Universe (e.g. Brinchmann et al. 2004; Peng et al. 2010), as well as in the distant Universe out to $z \gtrsim 2$ (e.g. Noeske et al. 2007; Daddi et al. 2007; Santini et al. 2009; Kajisawa et al. 2010; Bauer et al. 2011; Whitaker et al. 2012). A growing number of studies are now investigating various aspects of this relation; e.g. the origin of its scatter or its morphological dependence (Wuyts et al. 2011; Salmi et al. 2012). An interesting implication from some detailed studies of the local SFR– M_* relation is the “independence” of this relation with environment; Peng et al. (2010) studied local star-forming galaxies drawn from SDSS to show an excellent agreement in the SFR– M_* sequence between low-density and high-density environments. They argue that the environment

does change the star-forming galaxy *fraction*, but that it has very little impact on the SFR– M_* relation of those galaxies that are star-forming (see also Wijesinghe et al. 2012).

An observational challenge is to test this universality of SFR– M_* relation in the distant Universe. Any local relation may not necessarily be applicable for distant galaxies at $z \gtrsim 1$, where the average star-formation activity is about an order of magnitude higher (e.g. Hopkins & Beacom 2006; see also Sobral et al. 2013; Stott et al. 2013). Unfortunately, constructing a large, uniformly selected galaxy sample at such high redshifts is still challenging, which prohibits us from understanding the environmental impacts on the SFR– M_* relation in the distant Universe. Some earlier works have attempted to identify the environmental dependence (or its absence) of the SFR– M_* relation out to $z = 1$ (Vulcani et al. 2010; Li et al. 2011; McGee et al. 2011; Tyler et al. 2011; Muzzin et al. 2012), or out to $z = 2$ (Tanaka et al. 2010; Tanaka et al. 2011; Grützbauch et al. 2011; Koyama et al. 2013), but a full consensus on the environmental impacts on the SFR versus M_* relation has not yet been reached, because of the different sample selection and/or different environment definitions. Therefore, the second goal of this paper is to test the environmental dependence of the SFR– M_* relation out to $z \sim 2$ for the first time based on the purely $H\alpha$ -selected star-forming galaxy samples established in our recent two narrow-band $H\alpha$ survey projects; MAHALO-Subaru for clusters and HiZELS for general fields (see details in § 2.2).

The structure of this paper is as follows. In § 2, we present the wide-field Subaru and *Spitzer*/MIPS $24\mu\text{m}$ data of Cl0939, as well as the $H\alpha$ emitter samples at different redshifts and in different environments from our previous studies. We also describe how we estimate fundamental physical quantities such as stellar masses and star formation rates of galaxies. The main results and discussion of the paper are described in § 3. We first discuss MIR properties of $H\alpha$ emitters around Cl0939, particularly focusing on the nature of the red $H\alpha$ sources (§ 3.1–§ 3.2), and then we discuss the environmental dependence of the SFR versus M_* relation since $z \sim 2$ using all the $H\alpha$ selected galaxies (§ 3.3–§ 3.6). Finally, our conclusions are given in § 4. Throughout the paper, we adopt the standard cosmology with $\Omega_M = 0.3$, $\Omega_\Lambda = 0.7$, and $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, and we assume a Salpeter (1955) initial mass function (IMF). Magnitudes are all given in the AB system.

2 DATA AND ANALYSIS

2.1 The Cl0939 cluster

2.1.1 Subaru data

The Cl0939 cluster at $z = 0.41$ is one of the best-studied clusters at intermediate redshifts (e.g. Dressler & Gunn 1992; Dressler et al. 1994; Stanford, Eisenhardt, & Dickinson 1995; Smail et al. 1999; Sato & Martin 2006; Dressler et al. 2009; Oemler et al. 2009). In addition to these studies which focus on cluster central regions, very wide-field ($\sim 30' \times 30'$) optical broad-band and narrow-band imaging surveys of this cluster have been made by Kodama et al. (2001) and Koyama et al. (2011) using Suprime-Cam (Miyazaki et al. 2002) on the Subaru Telescope (Iye et al. 2004). The details of these Subaru data have already been presented in our previous papers (Kodama et al. 2001; Koyama et al. 2011), and so here we provide only a quick summary of the data. The broad-band data are analysed by Kodama et al. (2001). They discovered a 10 Mpc-scale filamentary large-scale structure around the

cluster based on the photometric redshift (photo- z) technique. The $H\alpha$ emitter search of this field is made by Koyama et al. (2011), using the narrow-band filter NB921 ($\lambda_c=9180\text{\AA}$) on Suprime-Cam, who identified >400 $H\alpha$ emitting galaxies around the cluster. In this paper, we use both the $H\alpha$ emitter and the photo- z selected cluster member catalogues presented in these studies. The photometry of the sources is performed with the SExtractor software (Bertin & Arnouts 1996). We primarily use MAG_AUTO as the total magnitudes for measuring physical quantities such as M_* or SFR, while we use $3''$ aperture (corresponding to 16 kpc) photometry for measuring galaxy colours.

2.1.2 *Spitzer MIPS 24 μ m data*

We retrieve the wide-field MIPS (Rieke et al. 2004) $24\mu\text{m}$ scan data of the C10939 field from the *Spitzer* Science Archive. The data covers a large part of our Suprime-Cam field of view (from cluster core to surrounding groups), so are well-suited for studying dust-obscured star formation activity around the cluster. The data was reduced from the Basic Calibrated Data stage (provided by the *Spitzer* Science Centre) using the MOPEX software following the procedure outlined in Geach et al. (2006). Source extraction was performed using SExtractor (Bertin & Arnouts 1996), with the criteria that a source consist of at least three contiguous pixels (each pixel is $2.5''$ square) at $>2\sigma$ above the background. We measure $16''$ diameter aperture fluxes, corresponding to ~ 3 times the FWHM of the point-spread function (PSF; FWHM of $5''$ at $24\mu\text{m}$). Using a curve-of-growth analysis on bright isolated point sources, and our completeness simulations, we find that $16''$ apertures recover ~ 75 per cent of the total flux, and we therefore correct the resulting fluxes by a factor of 1.33 to yield the total $24\mu\text{m}$ fluxes.

The $24\mu\text{m}$ catalogue contains 886 sources down to $200\mu\text{Jy}$ ($\sim 5\sigma$ limit) within the Suprime-Cam field. We perform cross-matching between the $24\mu\text{m}$ sources and our optical sources to construct a MIPS-detected cluster member catalogue. We search for optical sources within $3''$ radius from each $24\mu\text{m}$ source, and we select the nearest source as its optical counterpart. As a result, we find 162 MIPS sources which are likely to be associated with cluster member galaxies (photo- z members or $H\alpha$ emitters), although due to the poor spatial resolution of the MIPS $24\mu\text{m}$ data, it is sometimes difficult to identify the counterparts correctly. In order to avoid such ambiguous detections, we check all the sources by eye and exclude some heavily blended sources. Only 10 sources (6%) are excluded in this process (and they seem to be just normal galaxies randomly distributed on sky), so the exclusion of these sources does not affect our results at all. Overall, the MIPS-detected member catalogue contains total 152 sources, among which 33 sources are $H\alpha$ emitters. The relatively small number of the $H\alpha$ detected sources is not surprising, because the photo- z selected members include galaxies over a relatively broad range in redshift (we apply $0.30 \leq z_{\text{phot}} \leq 0.45$ for photo- z member criteria; see Koyama et al. 2011), while the $H\alpha$ emitters are considered to be secure cluster members located within the narrow redshift slice at the cluster's redshift ($0.39 \leq z \leq 0.41$).

2.1.3 *MIPS stacking analysis*

We note that the limiting flux of the $24\mu\text{m}$ data corresponds to $\text{SFR} \sim 8M_\odot \text{yr}^{-1}$, while our $H\alpha$ survey reaches down to $\text{SFR} \ll 1M_\odot \text{yr}^{-1}$ (without dust extinction correction). Indeed, the number of $H\alpha$ emitters which are individually detected at $24\mu\text{m}$ is

not large, because of the limited depths of the $24\mu\text{m}$ data. We therefore apply a $24\mu\text{m}$ stacking analysis for the $H\alpha$ emitters around C10939 to study the general properties of the fainter sources, by dividing the sample into subsamples selected by colour or environment (see § 3.2, § 3.3). We exploit median stacking at the positions of $H\alpha$ emitters down to an dust-uncorrected SFR of $0.25M_\odot \text{yr}^{-1}$ (which corresponds to a 5σ detection in $H\alpha$). Although our conclusion does not change even if we apply an average stacking, median stacking is preferred because it can minimize the effects from exceptionally luminous sources, as well as those from some luminous nearby sources.

The photometry on the MIPS stacked image is performed in the same way as for individual sources; i.e. $16''$ aperture photometry with the aperture correction of $\times 1.33$ (§ 2.1.2). Also, we apply bootstrap re-sampling (500 times) to obtain the $1-\sigma$ error-bars for the median $24\mu\text{m}$ fluxes. Note that we include all the $H\alpha$ emitters for the stacking, regardless of their individual $24\mu\text{m}$ detection, while we do not use photo- z sources (without $H\alpha$ detection) in order to avoid contaminant galaxies. We derive the total infrared luminosity (L_{IR}) from the stacked $24\mu\text{m}$ photometry using the Chary & Elbaz (2001) SED templates, and then compute the SFR using the combined $H\alpha$ and L_{IR} approach suggested by Kennicutt et al. (2009); $\text{SFR} = 7.9 \times 10^{-42} [L(H\alpha)_{\text{obs}} + 0.0024 \times L_{\text{TIR}}]$ (erg s^{-1}). This equation is derived from a tight correlation between the combined $H\alpha$ and total IR luminosities and the attenuation corrected SFR (see Kennicutt et al. 2009). We note that our conclusion does not change even if we calculate SFRs directly from the total IR luminosity using the Kennicutt (1998) relation (in this case we tend to derive SFRs which are higher by a factor of $\lesssim 2$), but the combined $H\alpha$ and IR approach should be more reliable, considering the moderate levels of star-formation activity (with $\lesssim 10M_\odot \text{yr}^{-1}$) of our galaxy sample.

2.2 $H\alpha$ emitter samples from our previous studies

2.2.1 *Cluster galaxy sample from MAHALO-Subaru*

In addition to the C10939 cluster ($z = 0.4$), we also use our similar, $H\alpha$ selected galaxies in a $z = 0.8$ cluster (RXJ1716+6708) and in a $z = 2.2$ proto-cluster (PKS1138–262), from our previous studies (see Koyama et al. 2010; 2013 for detailed descriptions of these data and the selection of $H\alpha$ emitters). The $H\alpha$ emitter samples are constructed with narrow-band surveys using NB119 ($\lambda_c = 1.19\mu\text{m}$ for $z = 0.8$) and NB2071 ($\lambda_c = 2.07\mu\text{m}$ for $z = 2.2$) on MOIRCS/Subaru (Ichikawa et al. 2006; Suzuki et al. 2008), as a part of the *MApping H-Alpha and Lines of Oxygen with Subaru* project (MAHALO-Subaru; see overview by Kodama et al. 2013). We use these cluster galaxy samples to study the redshift evolution of the SFR- M_* relation in cluster environments (see § 3.4). We note that our advantage is a perfect matching of their redshifts with our similar $H\alpha$ -selected control field galaxy samples from HiZELS (see below). Using only one cluster at each redshift might be too simplistic (e.g. our proto-cluster at $z = 2.2$ may not necessarily be the progenitor of our $z = 0.4$ and 0.8 clusters), but in this pioneering study, we assume that these $H\alpha$ emitter samples in these three cluster fields represent typical star forming galaxies in high-density environments at each redshift.

2.2.2 *Field galaxy sample from HiZELS*

In order to test the environmental dependence of the SFR versus M_* relation, we also need a control field galaxy sample. The High-

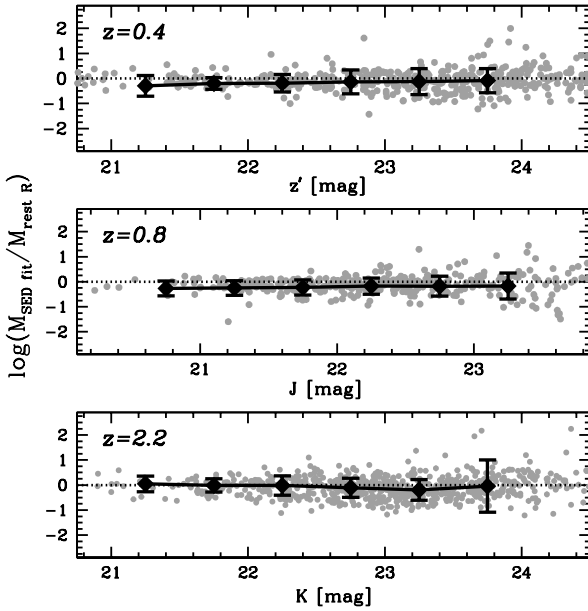


Figure 1. A comparison between the stellar mass estimates for HiZELS H α emitters derived from the rest-frame R -band magnitudes with one-colour correction (see text) and those from full SED fitting derived by Sobral et al. (in prep). The line-connected black points show the running median (and its associated scatter). It is clear that these two measurements agree with each other reasonably well over a wide luminosity range.

Z Emission Line Survey (HiZELS; Best et al. 2010; Sobral et al. 2013) is ideally suited for this purpose. This is not only because HiZELS is currently the largest narrow-band H α survey ever published, but also because three of their four targeted redshifts ($z = 0.4/0.8/2.2$) are perfectly matched with our cluster samples, allowing a direct cluster–field comparison based on the purely H α -selected galaxies. The HiZELS H α emitter samples are selected from the UDS and COSMOS fields (total ~ 2 deg 2 ; see also Geach et al. 2008; Sobral et al. 2009) by wide-field narrow-band imaging observations with NB921 on Subaru ($z = 0.4$), NB $_J$ on UKIRT ($z = 0.8$), and the NB $_K$ /H $_2$ filters on UKIRT/VLT ($z = 2.2$; see also Geach et al. 2012). We select H α emitters in exactly the same way as described in Sobral et al. (2013). We note that the selection and photometry of the HiZELS sources has been made with $3''$ and $2''$ aperture for $z = 0.4$ and $z = 0.8/2.2$ galaxies (see Sobral et al. 2013), respectively, while physical quantities of our cluster galaxy samples are measured with total magnitudes (MAG_AUTO from SExtractor). Therefore, we apply an aperture correction for the HiZELS sample based on the median difference between aperture magnitudes and total magnitudes for each redshift slice, although these corrections are negligibly small in our discussion (0.2 dex at maximum).

2.3 Stellar mass and H α -based star formation rate

In this subsection, we derive stellar masses (M_*) and star formation rates (SFR) of the H α emitters. The stellar masses of galaxies are ideally derived using a SED fitting approach including rest-frame near-infrared (NIR) photometry. However, the rest-frame NIR photometry is not available for our cluster galaxy samples. We therefore decide to estimate the stellar masses of galaxies based on their rest-frame R -band magnitudes. The conversions from observed magnitudes to M_* are determined using the model

galaxies developed by Kodama, Bell, & Bower (1999) (see also Kodama & Arimoto 1997; Kodama et al. 1998), and they are expressed by the following equations:

$$\log(M_*/10^{11} M_\odot)_{z=0.4} = -0.4(z' - 20.07) + \Delta \log M_{0.4}, \quad (1)$$

$$\log(M_*/10^{11} M_\odot)_{z=0.8} = -0.4(J - 21.14) + \Delta \log M_{0.8}, \quad (2)$$

$$\log(M_*/10^{11} M_\odot)_{z=2.2} = -0.4(K_{(s)} - 22.24) + \Delta \log M_{2.2}. \quad (3)$$

The final term in each equation (ΔM) accounts for the colour dependence of the mass-to-light ratio (M/L) predicted by the same galaxy model, and they correspond to:

$$\Delta \log M_{0.4} = 0.054 - 3.81 \times \exp[-1.28 \times (B - z')], \quad (4)$$

$$\Delta \log M_{0.8} = 0.085 - 2.48 \times \exp[-1.29 \times (R - J)], \quad (5)$$

$$\Delta \log M_{2.2} = 0.030 - 1.50 \times \exp[-1.11 \times (z' - K_{(s)})]. \quad (6)$$

The stellar masses derived with this “one-colour method” could be less accurate compared to those from full SED fitting, but we find that there is no systematic difference between the M_* from this method and those from the SED-fitting approach using HiZELS H α emitter samples for which full SED information are available (see Fig. 1). The two measurements are consistent with each other over a wide luminosity range (with a ≈ 0.3 dex scatter), verifying that our one-colour method works reasonably well. In this paper, in order to make a fair comparison between cluster and field samples at different redshifts, we use the above one-colour method for all H α emitters (including MAHALO and HiZELS samples).

We then calculate SFRs of H α emitters based on their H α luminosities. We first correct for the contribution of [NII] lines to the total NB fluxes, using an empirical correlation between $\text{EW}_{\text{rest}}(\text{H}\alpha + [\text{NII}])$ and the [NII]/H α ratio as described in Sobral et al. (2013). This relation has a large intrinsic scatter, but this method is believed to be more realistic approach compared with the conventional constant [NII] correction. We also apply a dust extinction correction to the H α flux of individual galaxies based on their stellar mass, as shown by Garn & Best (2010). This extinction correction could also be uncertain, given the large intrinsic scatter of the $A_{\text{H}\alpha}$ – M_* relation for local galaxies (see Garn & Best 2010). However, it is one of the most reliable estimators of dust extinction, and also the relation is reported to be unchanged out to $z \sim 1.5$ (Sobral et al. 2012). We therefore apply the same extinction correction to all H α emitters. Finally, we compute the SFRs of galaxies using the Kennicutt (1998) relation; $\text{SFR}_{\text{H}\alpha}(M_\odot \text{ yr}^{-1}) = 7.9 \times 10^{-42} L_{\text{H}\alpha}(\text{erg s}^{-1})$. We note that the [NII] line and the dust extinction corrections could be major sources of uncertainty, but this is currently inevitable because it is impossible to measure [NII] contribution or dust extinction for individual galaxies. We stress that we adopt the same procedure for all our sample (cluster and field galaxies), so that we consider the relative cluster–field comparison should be robust (see related discussions in § 4.5).

The stellar masses and SFRs derived above may not be very accurate for galaxies hosting active galactic nuclei (AGN), as their continuum light and H α line fluxes are contaminated by AGNs. For the field H α emitters (HiZELS), Garn et al. (2010) and Sobral et al. (2013) carried out a detailed study on the contribution of AGNs into their H α emitters sample. By applying various techniques for AGN selection (e.g. X-ray, radio, or emission-line ratios), they find that the AGN fraction amongst their H α emitters is as large as 10% out

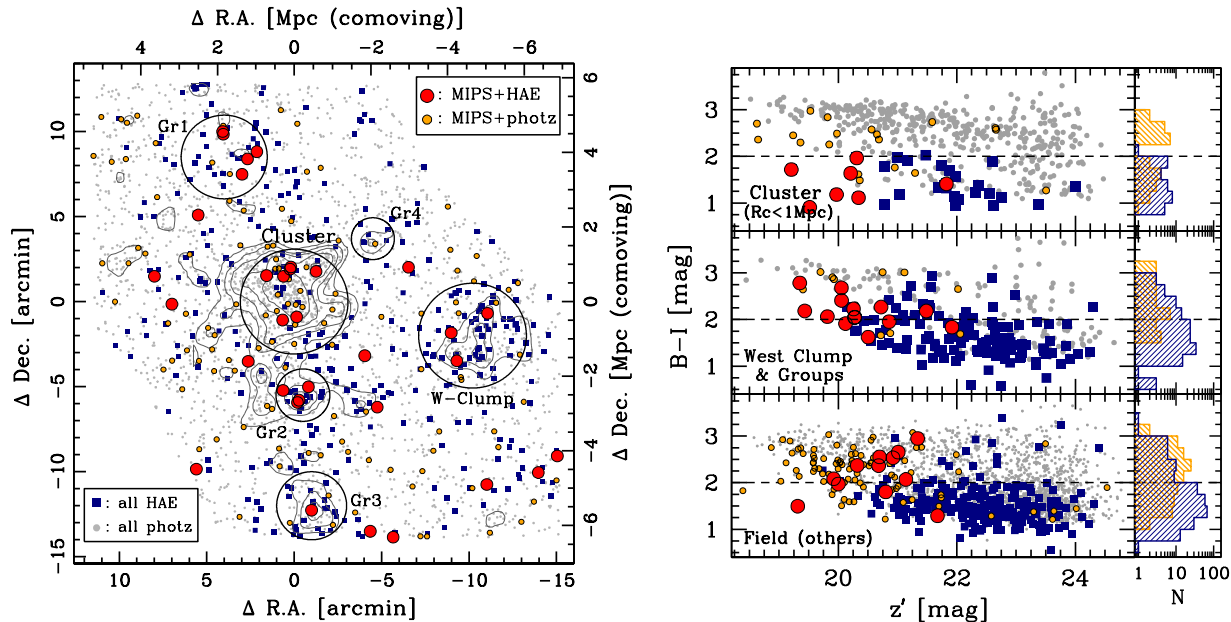


Figure 2. (Left): The 2-D map of galaxies around the cluster C10939 ($z = 0.41$). Galaxies within both Suprime-Cam and MIPS FoVs are shown. We plot $24\mu\text{m}$ -detected $\text{H}\alpha$ emitters, $24\mu\text{m}$ -detected photo- z members ($0.30 \leq z_{\text{phot}} \leq 0.45$), all $\text{H}\alpha$ emitters, and all photo- z members (the meanings of the symbols are shown in the plot). We also show the locations of the surrounding groups following the definitions in Koyama et al. (2011). Contours are drawn based on the surface number density of all member galaxies (same as fig. 5 of Koyama et al. 2011). (Right): The colour-magnitude diagram for each environment as defined in the left-hand panel. The meanings of the symbols are the same as the left-hand panel. The (blue) hatched and (orange) cross-hatched histograms show the colour distribution of $\text{H}\alpha$ emitters and $24\mu\text{m}$ sources, respectively.

to $z \sim 1$ or $\sim 15\%$ at $z > 1$. For cluster samples, we confirmed a few (up to five) AGNs in each cluster using X-ray imaging data or spectroscopic information (see Koyama et al. 2008; 2011; 2013). Considering the size of the total galaxy sample we use in this paper (~ 100 galaxies per cluster), it is roughly estimated that the AGN fraction in our cluster galaxies is $\approx 5\text{--}10\%$. Unfortunately, it is not possible to fully quantify the contribution of AGNs in our cluster $\text{H}\alpha$ emitters in the same way as for HiZELS samples (due to the lack of multi-wavelength data). In order to do a fair comparison between cluster and field samples, we do not exclude AGNs from our cluster or field $\text{H}\alpha$ emitters in the following discussions, but we note again that the AGN contribution is always small ($\approx 10\%$) and should not strongly bias the results.

3 RESULTS AND DISCUSSION

3.1 Panoramic $\text{H}\alpha$ and MIR view of the C10939 cluster

In Fig. 2, we show the spatial distribution of the MIR-detected $\text{H}\alpha$ emitters and the MIR-detected photo- z member galaxies. We also show all the $\text{H}\alpha$ emitters and photo- z member galaxies, and the locations of the West Clump and the four surrounding groups following the definitions in Koyama et al. (2011). Note that we only show the galaxies located within the overlapped regions between our Subaru and MIPS FoVs. In the right-hand panel of Fig. 2, we show the colour-magnitude diagram for cluster (1 Mpc from the cluster centre), group (including the west clump and four groups), and field environments. These colour-magnitude diagrams show an overall trend that the MIPS-detected sources tend to be luminous ($z' \lesssim 21$ mag) and to have red optical colours ($B-I \gtrsim 2$ mag).

We compare the $B-I$ colour distribution of the $\text{H}\alpha$ emitters in each environment (see histograms in Fig. 2). The lack of the red $\text{H}\alpha$ emitters ($B-I > 2$ mag) is clearly visible in the cluster environment,

while a large number of $\text{H}\alpha$ emitters tend to have red colours in the group environment. The Kolmogorov-Smirnov (KS) test shows that the probability that the colour distributions of $\text{H}\alpha$ emitters in clusters and groups environment are from the same parent population is 1%. The trend becomes less significant if we compare the group galaxies with field galaxies, but the KS test still suggests that the group and field galaxies are unlikely to be drawn from the same parent population (2.5%). This probability goes down to $< 2\%$ if we use “all” $\text{H}\alpha$ emitters sample within the Suprime-Cam FoV (including the galaxies located outside the MIPS data coverage). Furthermore, we calculate the fraction of the red galaxies with any indication of star formation activity (i.e. $\text{H}\alpha$ or MIR detection): $f = N_{\text{redSF}} / N_{\text{allred}}$. We find fractions of $4 \pm 1\%$, $17 \pm 3\%$, $11 \pm 1\%$ for cluster, group, and field environments, respectively. We therefore conclude that the red star-forming galaxies are most frequently seen in the group-scale environment (supporting our previous finding in Koyama et al. 2011); in other words, a non-negligible fraction of optically red galaxies in the group environment are still actively forming stars.

3.2 Stacked MIR properties as a function of galaxy colours

The colour-magnitude diagrams in Fig. 2 show that a fraction of the red $\text{H}\alpha$ emitters are individually detected at $24\mu\text{m}$, suggesting they are dusty sources rather than passive galaxies. However, the limited depth of the $24\mu\text{m}$ data prevents us from assessing the general MIR properties of the faint $\text{H}\alpha$ galaxies. We therefore apply a $24\mu\text{m}$ stacking analysis to study the properties of the $\text{H}\alpha$ emitters more generally. To do this, we divide the full $\text{H}\alpha$ emitter sample in the C10939 field into three colour bins (at $B-I = 1.5$ and 2.0), and perform stacking analysis as described in § 2.1.3. The results are shown in Fig. 3a. It is clear that the redder $\text{H}\alpha$ sources have higher SFRs than bluer sources. We also show in Fig. 3a the SFRs derived

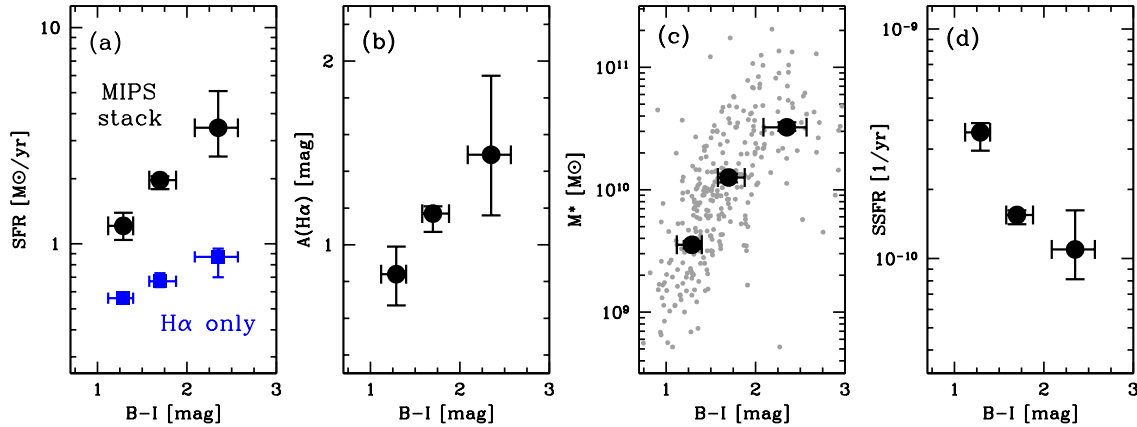


Figure 3. The stacked properties of the $H\alpha$ emitters as a function of $B-I$ colours (rest-frame $U-V$). The vertical error-bars show the 1σ uncertainty derived from the bootstrap re-sampling approach during the MIPS stacking (see text), while the horizontal error-bars show the 25–75% colour distribution of each subsample. (a) The SFRs derived from $24\mu\text{m}$ stacking analysis (black symbols). The blue symbols indicate the median values of $\text{SFR}_{H\alpha}$ (without dust extinction correction) for each subsample. (b) The $A_{H\alpha}$ value for each subsample, calculated from the ratio of SFR_{IR} and $\text{SFR}_{H\alpha}$. (c) The median stellar mass for each colour subsample (black), as well as those for individual sources (grey dots). (d) The SSFR for each subsample derived as SFR normalized by stellar mass. These four plots demonstrate that the red $H\alpha$ emitters are massive star-forming galaxies with higher SFR and higher dust extinction compared with normal blue $H\alpha$ emitters.

from the median $H\alpha$ flux for each subsample, without dust extinction correction. The same trend is still visible, but it is weaker than that obtained from MIPS stacking analysis. This result suggests that the redder $H\alpha$ galaxies are dustier than bluer galaxies. In Fig. 3b, we show the extinction at $H\alpha$ derived from a ratio of the total SFR (from MIPS stacking) and the dust-free $H\alpha$ SFR as a function of galaxy colours. This plot clearly shows that the red $H\alpha$ emitters have much higher extinction (with $A_{H\alpha} \sim 1.5$ mag) compared with the blue emitters ($A_{H\alpha} \lesssim 1$ mag), suggesting that the red $H\alpha$ emitters are dusty, star-forming galaxies.

We should note that the clear trend that the redder $H\alpha$ emitters have higher SFRs could be produced by a stellar mass difference between red and blue $H\alpha$ emitters. In Fig. 3c, we show the estimated M_* for the individual $H\alpha$ emitters (see § 2.3), as well as the median stellar mass for each colour subsample. The red $H\alpha$ emitters tend to have much higher stellar masses (by a factor of ~ 10) than the blue $H\alpha$ emitters, which is more significant than the difference in SFR (i.e. Fig. 3a). Consequently, it turns out that the red $H\alpha$ emitters tend to have lower specific star formation rate (SSFR) than blue $H\alpha$ galaxies (see Fig. 3d). However, the SSFR for the red $H\alpha$ emitters are still at the $\sim 10^{-10} \text{ yr}^{-1}$ level, which is about one order of magnitude higher than the usual definitions of “passive” galaxies (10^{-11} yr^{-1}), suggesting they are part of the “star-forming” population.

As a further check, we show in Fig. 4 (left) the results of our stacking results for the red and blue $H\alpha$ emitters on the $\text{SFR}-M_*$ diagram (we divide the red and blue $H\alpha$ emitter sample into two equal-sized stellar mass bins). Our data points show an excellent agreement with the best-fitted $\text{SFR}-M_*$ relation for the $z = 0.4$ $H\alpha$ emitters from HiZELS (Sobral et al. 2013), further supporting our conclusion that the red $H\alpha$ emitters are dusty star-forming galaxies, rather than dust-free passive galaxies. We recall that when deriving $\text{SFR}-M_*$ relation for the HiZELS sample, we adopt an M_* -dependent dust extinction correction to HiZELS data (see § 2.3). The excellent agreement (over the wide M_* range) between the SFRs derived from the MIPS stacking analysis and those from the independent, extinction corrected $H\alpha$ fluxes also supports the va-

lidity of our procedure for the dust extinction correction applied in this study.

3.3 Environmental dependence

We have shown that our $H\alpha$ emitters tend to be located on the general $\text{SFR}-M_*$ sequence (Fig. 4-left). An interesting question here is: does environment influence the $\text{SFR}-M_*$ relation? To answer this question, we attempt a similar analysis to § 3.2 but dividing the whole $z = 0.4$ $H\alpha$ emitter sample into four environment bins based on the local galaxy density (at $\log \Sigma_{10} = 1.6, 1.9, 2.2$). The density is calculated using all cluster member galaxies (photo- z selected and $H\alpha$ selected) with the nearest-neighbour approach, calculated within a radius to the 10th-nearest neighbour from each source. We further divide each environment subsample into two equal-sized stellar mass bins, and perform the $24\mu\text{m}$ stacking analysis. The results are shown in Fig. 4 (right). The stacking analysis becomes challenging, particularly for the lower-mass sources due to the limited sample size, but broadly speaking, all our data points are likely to be located on the same $\text{SFR}-M_*$ sequence. This suggests that the $\text{SFR}-M_*$ (or $\text{SSFR}-M_*$) relation for star-forming galaxies does not strongly correlate with the environment at $z = 0.4$.

Using the same environmental subsamples, we also test the SFR -density relation for star-forming galaxies at $z = 0.4$. In Fig. 5 (left), we show the SFRs from $24\mu\text{m}$ stacking analysis as a function of galaxy density. It is interesting to note that the SFR *increases* with increasing galaxy density (by a factor of $\sim 3-4$), and the trend is confirmed for both red and blue $H\alpha$ emitter samples. In contrast, this trend becomes much less significant when we normalise the SFRs by M_* to compare the specific SFR (see Fig. 5 right). The fact that the SSFR shows a weaker environmental trend than SFR (Fig. 5) implies that the SFR excess detected in the high-density environment would largely be explained by the M_* difference between the different environments. We note, however, that the M_* distribution amongst $H\alpha$ emitters does not seem to be strongly dependent on environment (as shown by histograms in Fig. 4b). We can still find a weak trend that the galaxies in high-density environment tend to be more massive, but the difference is at the ~ 0.2 dex

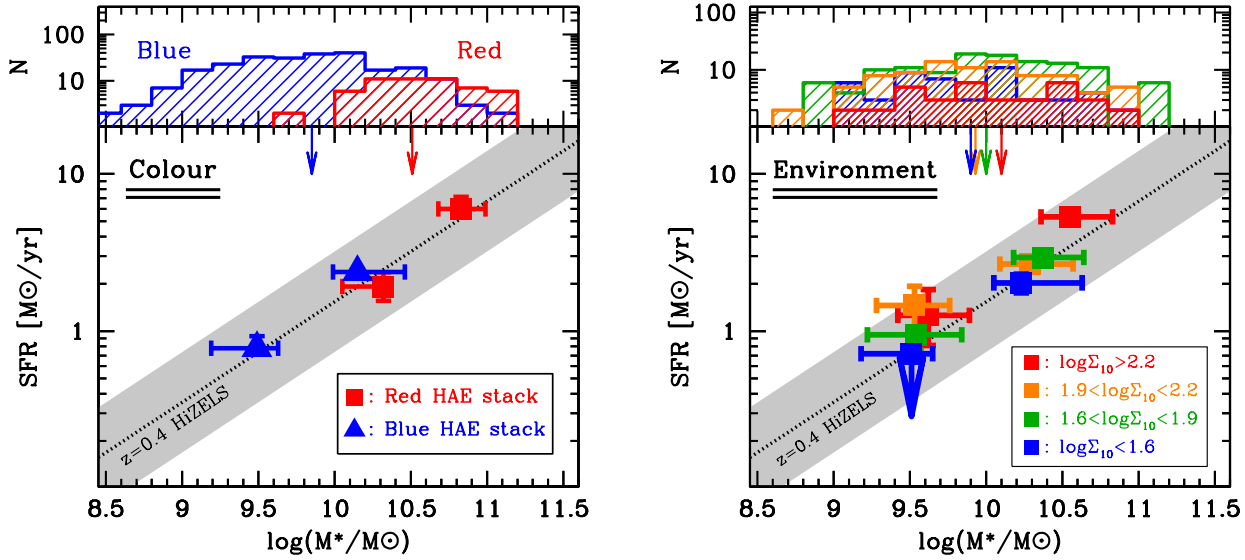


Figure 4. (Left): The median SFRs from $24\mu\text{m}$ stacking analysis for the red and blue $\text{H}\alpha$ emitters as a function of stellar mass. The red and blue samples are further divided into two equal-sized stellar mass bins. The dotted line and the shaded region show the best-fitted SFR- M_* relation and its scatter for the HiZELS $z = 0.4$ $\text{H}\alpha$ emitter sample from Sobral et al. (2013). All our data points show an excellent agreement with the HiZELS relation. (Right): The same plot as the left-hand panel, but for our environmental subsamples (each sample is further divided into two equal-sized stellar-mass bins). In both panels, we show the M_* distribution (histograms) and the median stellar masses (arrows) for each sample. In each plot, the vertical error-bars are from the bootstrap re-sampling in the MIPS stacking analysis, while the horizontal error-bars show the 25–75% M_* distribution for each subsample.

level at maximum. The weakness or lack of the environmental dependence of the stellar mass distribution amongst star-forming galaxies is consistent with some recent studies (Greene et al. 2012; Giodini et al. 2012), but this small M_* difference may not be able to fully account for the significant SFR increase towards high-density environment. We therefore speculate that the SFR excess in the high-density regions could be explained by a “mixed effect” of both slightly higher stellar masses *and* a small SSFR excess (both at ~ 0.2 dex level) in the high-density environment. Indeed, the small SSFR excess is visible as the remaining positive slope in the SSFR versus Σ_{10} plot (see Fig. 5 right), which is also equivalent to the small positive offset of the stacked data points for high-density environments in the SFR- M_* plot (Fig. 4 right). However, a firm conclusion will require future studies with larger galaxy samples, and so our reasonable conclusion at this stage is that the overall SFR- M_* relation is not strongly correlated with the environment at $z = 0.4$.

It should be noted again that the star-forming activity of galaxies presented here only focuses on star-forming galaxies, and that we do not include passive galaxies in our analysis. Indeed, we showed in Koyama et al. (2011) that the $\text{H}\alpha$ emitter fraction is a strong function of environment, showing a significant decline towards the cluster core. The important message from our current analysis is that the SFRs of the uniformly $\text{H}\alpha$ -selected galaxies depend on environment in the sense that galaxies in high-density environment have higher SFRs. This would suggest that the star-forming galaxies “surviving” in the very high-density environment have experienced an accelerated stellar mass growth in their past (e.g. through starburst activity), while it is also possible that they have been forming stars at constant rates over a longer period of time than general field galaxies (i.e. nature effects). Unfortunately, we cannot discuss the past of their star formation history with our photometric data alone. In a future paper, we will investigate their

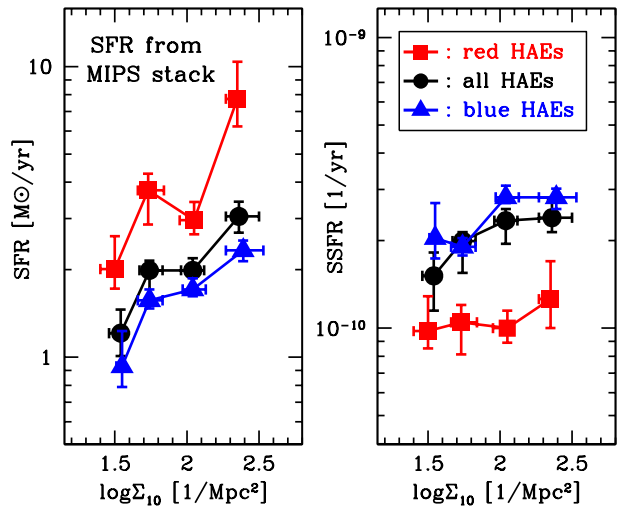


Figure 5. The SFR (left) and SSFR (right) for our $\text{H}\alpha$ emitters derived from $24\mu\text{m}$ stacking as a function of the local galaxy density (Σ_{10}). The density is calculated with all member galaxies (photo- z members and $\text{H}\alpha$ emitters), while we include only $\text{H}\alpha$ emitters in the stacking analysis. The SFR increases with increasing density regardless of their colours, while the trend becomes much flatter if we compare SSFR. The vertical error-bars show the $1-\sigma$ distribution from the bootstrap resampling and the horizontal error-bars show the 25–75% distribution, same as previous figures.

spectroscopic properties through an intensive spectroscopic follow up programme on this Cl 0939 field (Nakata et al. in preparation).

3.4 The evolving SFR- M_* relation since $z \sim 2$

We have examined the star-forming activity of galaxies and its dependence on stellar mass and environment at $z = 0.4$. Here, we dis-

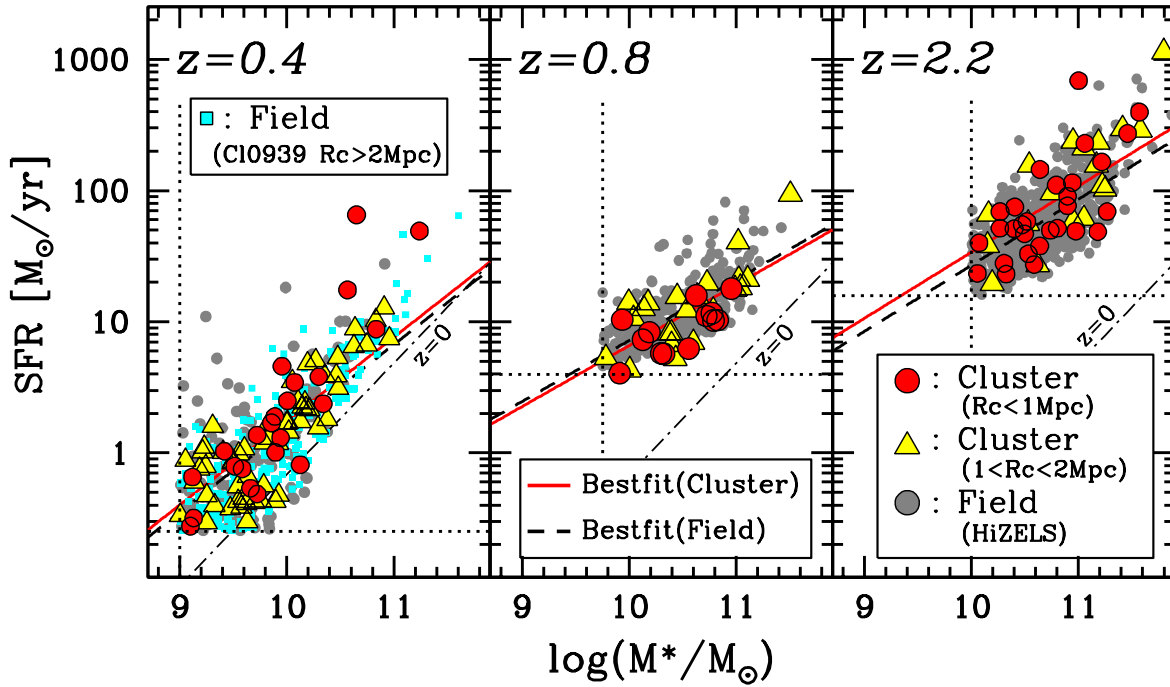


Figure 6. The SFR– M_* relation for the H α -selected galaxies at $z = 0.4, 0.8, 2.2$. In each panel, we plot the H α emitters in cluster environment located within $R_c < 1$ Mpc and at $1 < R_c < 2$ Mpc from the cluster centre. We also plot the H α emitters in general field environment selected from HiZELS at the same redshifts. Note that all the H α emitters plotted here are selected as those having $\text{EW}(\text{H}\alpha + [\text{NII}])_{\text{rest}} > 30 \text{ \AA}$. The dotted lines show the SFR and M_* cut applied at each redshift. The dashed line shows the best-fitted SFR– M_* relation for the HiZELS sample, while the solid line shows the relation for the cluster sample assuming the same slope as the HiZELS relation. The dot-dashed line is the local ($z = 0$) relation derived from the equation provided by Whitaker et al. (2012). It is clear that the SFR– M_* relation evolves significantly since $z \sim 2$ in both cluster and field environment, while at fixed redshifts, the environmental dependence of the SFR– M_* relation seems to be very small.

cuss the environmental dependence of the evolution of star-forming galaxies in a broader context, particularly focusing on the evolution and environmental dependence of the SFR versus M_* relation across cosmic time. By compiling all of the H α emitter samples together (including our MAHALO and HiZELS samples; see § 2.2), we examine the SFR versus M_* sequence in cluster and field environments at $z = 0.4, 0.8$, and 2.2 . We note that, while all the samples are selected based on the H α line, the EW cut applied in each survey is slightly different. Therefore we decide to apply the same (rest-frame) EW cut to all the sample, $\text{EW}_{\text{rest}}(\text{H}\alpha + [\text{NII}]) = 30 \text{ \AA}$, down to which all our H α data are complete. We also note that it is not possible to make a strictly fair comparison between clusters at different redshifts; for example, a density-based definition of environment requires accurate membership determination, while using the virial radius could be misleading because our high- z clusters are not virialised yet. Therefore, we instead use galaxies within 2 Mpc (in physical scale) from each cluster centre as “cluster” galaxies in the following discussion.

In Fig. 6, we show all the H α -selected galaxies in clusters and field environments at each redshift. We use different symbols for the cluster H α emitters located within $R_c < 1$ Mpc, $1 < R_c < 2$ Mpc, and the field H α emitters from HiZELS (see labels in the plot). For $z = 0.4$, we also show the H α emitters located at $R_c > 2$ Mpc from the CI0939 cluster. This plot clearly shows that the SFR– M_* relation evolves with redshift, while the relation is always independent of environment out to $z \sim 2$, qualitatively consistent with the situation in the local Universe (e.g. Peng et al. 2010; Wijesinghe et al. 2012). We also show the best-fitted SFR– M_* relation in each panel. It can be seen that the offset between the relation for cluster and field galaxies is always small. The slope of the SFR– M_* relation

tends to be steeper for the lower-redshift samples, which is also qualitatively consistent with previous works (but we note that the best-fit relation drawn on the plot is uncertain because our star-forming galaxy samples are not completely stellar-mass limited).

We here examine the galaxy distribution on the SFR– M_* plane more in detail. In Fig. 7, we show the distributions of M_* , SFR, and the offset from the main sequence (best-fitted SFR– M_* for field galaxies). The shaded histograms are for field H α emitters, and the hatched histograms are for cluster H α emitters. Broadly speaking, it seems that there is no significant environmental difference at any of the three epochs. For $z = 0.4/0.8$, the KS test suggests that it is unlikely that the cluster and field galaxies are from a different parent population, while for $z = 2.2$ we find a possible trend that cluster galaxies have a small excess in all three properties (KS test shows $< 1\%$ probability that the cluster and field samples are from the same parent population). We note that this trend is qualitatively consistent with some earlier studies showing a higher stellar mass in star-forming galaxies in $z > 2$ proto-clusters (e.g. Steidel et al. 2005; Kurk et al. 2009; Hatch et al. 2011; Matsuda et al. 2011; Koyama et al. 2013). Our current analysis also supports the idea that the star-forming galaxies in proto-cluster environment tend to be more massive than the general field galaxies, and this M_* excess would also account for the SFR excess in the proto-cluster environment; this may represent an accelerated galaxy growth in the early phase of the cluster assembly history.

Finally, we quantify the evolution of the star-forming activity of star-forming galaxies in the cluster and field environment. In Fig. 8, we plot the redshift evolution of the SSFR of H α emitters in clusters (red squares) and in field environment (black squares) at

the stellar mass of $\log(M_*/M_\odot) = 10$. It is found that the SSFR of $H\alpha$ -selected galaxies evolves significantly, going approximately as $(1+z)^3$, since $z \sim 2$ in both clusters and the field. This decline of SSFR is in good agreement with many studies of the cosmic star formation history (e.g. Yoshida et al. 2006; Karim et al. 2011). Therefore, an important indication from this study is that the evolution of star-forming galaxies in cluster environments just follows the same evolutionary track as that of general field galaxies.

3.5 Comparison with other studies

Studying the environmental dependence of galaxy star formation activity in the distant Universe is obviously an important step towards understanding the physical processes which drive the environmental effects. Since discussion on the “reversal” of the SFR–density relation was invoked by Elbaz et al. (2007), there has been much debate about the role of environment in the distant Universe. We note that our results are qualitatively consistent with some recent studies. For example, Sobral et al. (2011) used $H\alpha$ emitters sample at $z = 0.8$ selected from HiZELS to show higher median SFRs (by a factor of ~ 2 – 3) in high-density environment compared to low-density environment. They also showed that the stellar mass of $H\alpha$ emitters is weakly correlated with the environment (with ~ 0.3 dex increase in their highest-density bins), which also agrees with our finding. Tran et al. (2009) analysed MIR data of a “super-group” environment at $z = 0.37$ to show that the characteristic IR luminosity (L^*) in the group environment is higher than that in the field, based on the analysis of the IR luminosity function (see also Chung et al. 2010). Related to this, some studies of distant clusters show a peak of star formation activity at a certain galaxy density which corresponds to group or cluster outskirts environment (e.g. Poggianti et al. 2008; Koyama et al. 2010; Geach et al. 2011).

On the other hand, a growing number of studies recently have reported a weakness or absence of any relation between SSFR and environmental density amongst star-forming galaxies at least out to $z \sim 1$, or possibly to $z \sim 2$. In the local Universe, Balogh et al. (2004) showed that the EW($H\alpha$) distribution (equivalent to SSFR distribution) amongst star-forming galaxies is independent of environment. More recent studies also indicated that the SFR– M_* relation for local star-forming galaxies does not correlate with the environment (e.g. Peng et al. 2010; Wijesinghe et al. 2012). Similar suggestions have also been made for distant star-forming galaxies as well. For example, McGee et al. (2011) studied a large sample of $z = 0.4$ group galaxies to show that the average SSFRs of star-forming galaxies are the same in groups as in field environments. Muzzin et al. (2012) also showed that SSFR of star-forming galaxies is independent of environment at fixed stellar mass from their detailed spectroscopic survey of $z \sim 1$ cluster galaxies (see also Greene et al. 2012). Furthermore, our recent studies of distant (proto-)clusters also find a hint that the SSFR of star-forming galaxies is independent of environment at fixed stellar mass out to $z \sim 2.5$ (e.g. Tadaki et al. 2012; Hayashi et al. 2011; 2012; Koyama et al. 2013).

However, it should be noted that the independence of the SFR– M_* relation for star-forming galaxies with environment in the distant Universe is still controversial (Patel et al. 2011; Vulcani et al. 2010; Li et al. 2011). Indeed, this kind of analysis could be highly sensitive to the sample selection, the measurement of star formation rates, and the definitions of environment (see e.g. Patel et al. 2009). Our data are purely $H\alpha$ selected (for both cluster and field galaxies), and in this sense our cluster–field comparison would be robust. One possible bias is that our data are complete only for relatively

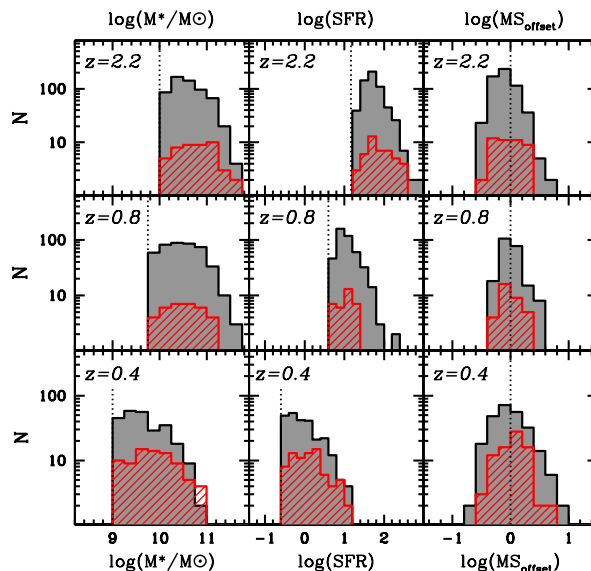


Figure 7. The distribution of M_* (left), SFR (middle), and offsets from the main sequence of field galaxies (right) at each redshift. The shaded histograms show the results for HiZELS sample, while the hatched histograms show the results for cluster ($R_c < 2$ Mpc) galaxies. The vertical dotted lines in the left and middle panels show the M_* or SFR cut we applied for each redshift sample, while the dotted lines in the right-hand panel show the location of the zero-offset. The actual difference between cluster and field galaxies is always small (~ 0.1 – 0.2 dex at maximum), but we note that a statistical test suggests that the two distributions may be different for our $z = 2.2$ sample in the sense that the cluster galaxies have higher M_* and higher SFR (see text).

strong emitters; we recall that our definition of star-forming galaxies is $EW > 30\text{\AA}$, so that we cannot discuss faint, low-EW sources. It is likely that such low EW sources do exist in both environments (and perhaps they may be more numerous in cluster environment). However, as reported in Sobral et al. (2011), such low-EW sources tend to be dominated by massive galaxies with relatively low SFR (i.e. largely “switched-off” population), so that we expect that it is unlikely that such low-EW sources would have a significant impact on our discussions on “star-forming” galaxies.

3.6 Interpretation and Caveats

Our main finding of this study is that the SFR versus M_* relation for the $H\alpha$ -selected galaxies does not strongly depend on the environment at any time in the history of the Universe since $z \sim 2$. This is a similar suggestion by Peng et al. (2010), who used COSMOS data to study the redshift evolution of the SSFR of (blue) star-forming galaxies, and find no environmental difference since $z \sim 1$. The independence of the SFR– M_* relation with environment could be explained if the environmental quenching is a rapid process (see e.g. Muzzin et al. 2012). That is, the environment instantly shuts down the star-formation activity of galaxies once the environmental effects are switched on, so that declined star-formation is not observed (because our galaxy samples are selected with $H\alpha$). Therefore a naive interpretation of our result would be that the major environment quenching mechanisms are always fast-acting in the history of the Universe since $z \sim 2$.

An important, but unexplored issue is the contribution of AGNs. While most of our $H\alpha$ -selected galaxies are likely to be

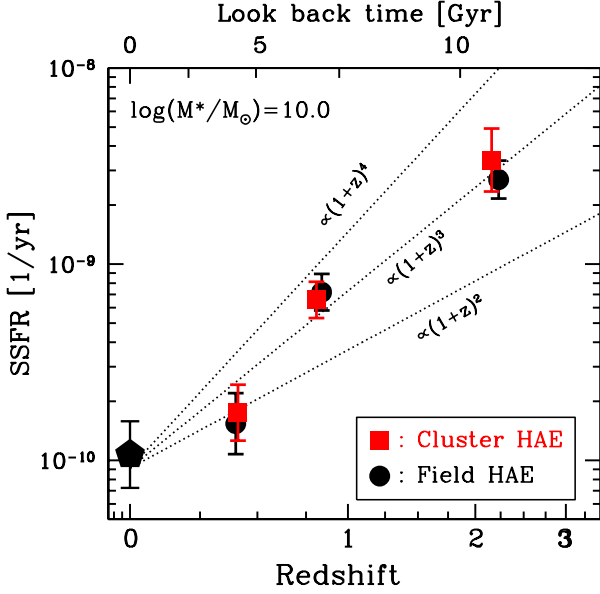


Figure 8. The redshift evolution of the SSFR at $M_* = 10^{10} M_\odot$ derived from the best-fitted SFR– M_* relation for cluster (red squares) and field (black circles) galaxies. The error-bars represent the standard deviation around their best-fit relation (see Fig. 6). The dotted lines are the evolutionary tracks following $\propto (1+z)^2$, $\propto (1+z)^3$, and $\propto (1+z)^4$, to guide the eye. The local data point is derived by adopting $z = 0$ in the equation of Whitaker et al. (2012).

powered by star formation (see § 2.3), there still remains a possibility that the AGN contribution could be dependent upon redshift, mass, and environment. For example, Popesso et al. (2011) carried out a detailed FIR study of the star forming activity of galaxies at $z \sim 1$ using *Herschel* data. They find that, while overall the SSFR– M_* relation does not depend on environment, the reversal of the SFR–density relation could be produced by very massive galaxy population. They also noted that the inclusion of AGNs into the analysis could also lead to an apparent reversal of the SFR–density relation. Therefore, more detailed studies of individual galaxies (including spectroscopy) are clearly needed to unveil the role of AGNs, as a future step of this study.

Another caveat on our result concerns the prediction of dust extinction correction. We applied the empirical correction based on the $A_{H\alpha}$ – M_* correlation established for local galaxies (see § 2.3), which has a large intrinsic scatter (Garn & Best 2010). The relation is reported to be unchanged out to $z \sim 1.5$ (Garn et al. 2010; Sobral et al. 2012), and as shown in Garn et al. (2010), the $A_{H\alpha}$ values (from the ratio of IR and $H\alpha$ luminosities) show no significant environmental variations. Indeed, as we showed in § 3.3 (Fig. 4), our results are unchanged even if we derive SFRs from a MIR stacking analysis, suggesting that this effect is not large. Using our C10939 sample presented in this paper, we further check the extinction values ($A_{H\alpha}$ from MIPS stacking analysis) for cluster ($R_c < 2$ Mpc) and field ($R_c > 2$ Mpc) galaxies. We find that the field galaxies have slightly (~ 0.2 mag) higher extinction, while we find an opposite trend if we use the density-coded definitions of environment (galaxies in high-density environment tend to be more extinguished). Therefore, it is not possible to say anything about the environmental dependence of the dust properties of galaxies from our data alone (due to limited sample size), but in any case, it is unlikely that the dust properties in star-forming galaxies are significantly different between different environments.

4 SUMMARY AND CONCLUSIONS

In this paper, we study the evolution and environmental dependence of the SFR– M_* correlation for star-forming galaxies since $z \sim 2$. We first present the MIR properties of the $H\alpha$ -selected star-forming galaxies in a rich cluster at $z = 0.4$ (C10939), and then we compare the $z = 0.4$ galaxies with our similar, $H\alpha$ -selected galaxies at different redshifts and in different environments. Our findings are summarized as follows:

(1) The red $H\alpha$ emitters, which are reported to be most frequently seen in the group-scale environment at $z = 0.4$, are dusty red galaxies rather than (semi-)passive galaxies. Using a wide-field *Spitzer*/MIPS $24\mu\text{m}$ dataset, we find that a large number of massive red $H\alpha$ sources are individually detected at $24\mu\text{m}$, suggesting they are luminous and dusty. Also, with a stacking analysis, we confirm a more general trend that the red $H\alpha$ sources tend to have higher SFRs with stronger dust extinction compared with normal blue $H\alpha$ emitters.

(2) We also find that the median SFR of $H\alpha$ emitters (derived from the MIR stacking analysis) increases with increasing galaxy number density at $z = 0.4$. This result is confirmed for both red and blue $H\alpha$ emitters, while the trend becomes much weaker if we compare their SSFR. We note that there still remains a positive correlation between SSFR and galaxy number density, and therefore we speculate that the SFR excess in the high-density environment can be caused by a mixed effect of both slightly higher M_* and a small SSFR excess (both at ~ 0.2 dex level) in high-density environment, implying that the M_* difference may not be able to fully account for the SFR increase in the high-density regions.

(3) Using our large $H\alpha$ -selected galaxy samples in distant cluster environments (from MAHALO-Subaru) and in general field environments (from HiZELS) at $z = 0.4, 0.8, 2.2$, we examine the environmental dependence of the SFR– M_* relation across cosmic time. We find that the SFR– M_* relation evolves with cosmic time, but that there is no detectable environmental variation in the SFR– M_* relation at any of these redshifts. This result suggests that, at fixed redshift, the overall star-forming activity amongst star-forming galaxies is not strongly correlated with environment once their stellar mass is fixed. On the other hand, for the $z = 2.2$ samples, we find a possible trend that the (proto-)cluster galaxies tend to have higher M_* (hence higher SFR), suggesting an accelerated galaxy mass growth in the early phase of cluster assembly.

(4) Based on the SFR– M_* relation we derived for cluster and field galaxies at $z = 0.4, 0.8, 2.2$, we also examine the evolution of the SSFR for star-forming galaxies (at the fixed mass of $M_* = 10^{10} M_\odot$). We find that the SSFR evolves significantly, as $(1+z)^3$, in both cluster and field environments. This means that the star-forming galaxy evolution in cluster environments just follows the same evolutionary track as that of field galaxies, which is most simply interpreted as implying that the primary physical driver of the environmental quenching is always a fast-acting process at any time in the history of the Universe since $z \sim 2$.

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